

SOLAR INDICES BULLETIN

OCTOBER 2005

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ISSN 1046-1914

♦ SOLAR RADIO EMISSIONS

The quiet Sun emits radio energy with a slowly varying intensity. These radio fluxes, which stem from atmospheric layers high in the chromosphere and low in the corona, change gradually from day-to-day, in response to the number and size of spot groups on the solar disk. The table below gives daily measurements of this slowly varying emission at selected wavelengths between about 1 and 100 centimeters. Many observatories record quiet-sun radio fluxes at the same local time each day and correct them to within a few percent for factors such as antenna gain, bursts in progress, atmospheric absorption, and sky background temperature. At 2800 megahertz (10.7 centimeters) flux observations summed over the Sun's disk have been made continuously since February 1947.

♦ SOLAR FLUX TABLE

Numbers in parentheses in the column headings below denote frequencies in megahertz. Each entry is given in solar flux units—a measure of energy received per unit time, per unit area, per unit frequency interval. One solar

flux unit equals $10^{-22} \text{ J/m}^2\text{Hzsec}$. During low periods of solar activity, the flux never falls to zero, because the Sun emits at all wavelengths with or without the presence of spots. The lowest daily Ottawa flux since 1947 occurred on November 3, 1954. On that day the observed noon value dropped to 62.6 units; the highest observed value of 457.0 occurred on April 7, 1947.

The preliminary observed and adjusted Penticton fluxes tabulated here are the "Series C" values reported by Canada's Dominion Radio Astrophysical Observatory in Penticton, British Columbia. Observed numbers are less refined, since they contain fluctuations as large as $\pm 7\%$ from the continuously changing sun-earth distance. Adjusted fluxes have this variation removed; they show the energy received at the mean distance between the Sun and Earth. Gaps in the Palehua, Hawaii (PALE), data reflect equipment problems. Fluxes measured either at Sagamore Hill, Massachusetts, or at San Vito, Italy, will be substituted for frequencies at which many Palehua values are missing.

OCTOBER 2005 PRELIMINARY SUNSPOT NUMBERS AND SOLAR RADIO FLUX

Day	Sunspot Number	Obs Flux Pentic (2800)	Solar Flux Adjusted to 1 Astronomical Unit								
			PALE (15400)	PALE (8800)	PALE (4995)	Pentic (2800)	PALE (2695)	PALE (1415)	PALE (610)	PALE (410)	PALE (245)
01	9	72	477	224	130	72	77	40	29	22	9
02	8	75	486	231	133	75	80	41	32	21	10
03	10	74	491	222	133	74	79	41	29	24	10
04	14	83	491	237	142	83	87	45	33	26	10
05	17	81	491	231	137	81	85	45	32	23	10
06	15	80	468	224	133	79	83	45	30	26	10
07	14	79	487	230	134	79	65	45	31	25	12
08	12	78	494	231	134	78	84	46	35	22	10
09	10	79	467	225	133	79	81	46	35	26	10
10	16	79	482	229	135	79	80	46	39	29	10
11	9	78	488	228	133	77	81	46	37	23	11
12	14	77	487	221	134	77	82	45	35	24	11
13	0	78	487	222	135	78	82	46	26	20	11
14	8	78	490	225	134	78	82	45	37	25	12
15	8	80	493	229	137	79	78	46	38	26	11
16	8	79	480	224	136	79	81	44	38	25	11
17	8	78	477	229	134	78	82	44	37	25	11
18	8	78	488	226	133	78	82	45	35	24	11
19	16	78				77					
20	9	77	479	224	130	76	79	42	36	22	10
21	8	75	480	225	128	75	78	42	32	22	10
22	7	75	459	211	127	74	75	41	33	21	10
23	8	74	479	223	129	73	76	40	33	21	10
24	0	73	485	225	127	73	75	40			
25	0	73	480	222	129	72	77	40	33	21	9
26	0	72	476	223	127	71	75	39	31	21	9
27	0	72	479	221	129	71	75	38	29	19	9
28	0	73	480	223	131	72		39	28	19	9
29	8	74	494	223	129	73	77	40	31	20	10
30	9	76	483	220	132	75		40	30	22	11
31	12	78	476	227	130	77	77	41	30	21	10
Mean	8	77	482	225	132	76	79	42	32	22	10

SEPT 2005 FINAL FLUX

Observed Pentic (2800)	Adjusted Pentic (2800)
79.2	80.6
77.1	78.5
74.2	75.4
74.6	75.8
75	76.2
83.4	84.7
92.2*	93.6*
94.1	95.5
99.2*	100.6*
116	117.6
109.7	111.1
118	119.5
113.6*	115.0*
116.6	117.9
119.4	120.6
111.8*	112.9*
103.9	104.9
102.2	103.2
91.1	91.9
87.8	88.5
86	86.7
83.7	84.3
82.8	83.3
81.4	81.9
81	81.4
*1700UT	Reading
81.3	81.7
76.9	77.2
74.6	74.9
73.8	74
72.2	72.4
91.1	92.1

♦ SUNSPOT COUNTS

In 1848 the Swiss astronomer Johann Rudolph Wolf introduced a daily measurement of sunspot number. His method, which is still used today, counts the total number of spots visible on the face of the Sun and the number of groups into which they cluster, because neither quantity alone satisfactorily measures the level of sunspot activity.

An observer computes a daily sunspot number by multiplying his estimated number of groups by ten and then adding this product to his total count of individual spots. Results, however, vary greatly, since the measurement strongly depends on observer interpretation and experience and on the stability of the Earth's atmosphere above the observing site. Moreover, the use of Earth as a platform from which to record these numbers contributes to their variability, too, because the Sun rotates and the evolving spot groups are distributed unevenly across solar longitudes. To compensate for these limitations, each daily international number is computed as a weighted average of measurements made from a network of

cooperating observatories. The international sunspot numbers tabulated on page 1 are provisional values taken from a bulletin prepared monthly by Pierre Cugnon of the SUNSPOT INDEX DATA CENTER, 3 avenue Circulaire, B-1180 BRUXELLES, BELGIUM. The October 2005 data combine observations from 45 stations. (<http://sidc.oma.be>)

♦ HISTORICAL SUNSPOT COUNTS

How do sunspot numbers in the table on page 1 compare to the largest values ever recorded? The highest daily count on record occurred December 24-25, 1957. On each of those days the sunspot number totaled 355. In contrast, during years near the spot cycle minimum, the count can fall to zero. Today, much more sophisticated measurements of solar activity are made routinely, but none has the link with the past that sunspot numbers have. Our archives, for example, include reconstructed daily values from January 8, 1818; monthly means from January 1749; and yearly means beginning in 1700.

SMOOTHED (OBSERVED AND PREDICTED) SUNSPOT NUMBERS: CYCLES 22 AND 23

Year	Jan	Feb	Mar	Apr	May	Jun	Jul	Aug	Sep	Oct	Nov	Dec	Mean
1993	71	69	67	64	60	56	55	52	48	45	41	38	56
1994	37	35	34	34	33	31	29	27	27	27	26	26	30
1995	24	23	22	21	19	18	17	15	13	12	11	11	17
1996	10	10	10	9	8*	9	8	8	8	9**	10	10	9
1997	10	11	14	17	18	20	23	25	28	32	35	39	23
1998	44	49	53	57	59	62	65	68	70	71	73	78	62
1999	83	85	84	86	91	93	94	98	103	108	111	111	96
2000	113	117	120	120.7#	119	119	120	119	116	115	113	112	117
2001	109	104	105	108	109	110	112	114	114	114	115	115	111
2002	114	115	113	111	109	106	103	99	95	91	85	82	102
2003	81	79	74	70	68	65	62	60	60	58	57	55	66
2004	52	49	47	46	44	42	40	39	38	36	35	34	42
2005	35	34	34	32	30	29	28	27	26	25	24	23	29
					(2)	(4)	(5)	(6)	(7)	(8)	(8)	(8)	(4)
2006	22	21	20	20	19	19	18	17	16	15	14	13	18
	(8)	(9)	(9)	(10)	(10)	(11)	(11)	(10)	(10)	(10)	(10)	(10)	(10)
2007	13	12	12	12	12	12	13	13	14	15	16	17	13
	(10)	(10)	(10)	(10)	(10)	(11)	(12)	(14)	(16)	(17)	(19)	(21)	(13)

*May 1996 marks Cycle 22's mathematical minimum. **October 1996 marks the consensus Cycle 22 minimum which NGDC is now using.

April 2000 marks Cycle 23 maximum.

♦ SUNSPOT NUMBER PREDICTIONS

For the end of Solar Cycle 22, and the beginning of Cycle 23, the table gives smoothed sunspot numbers up to the one calculated that first uses the most recently measured monthly mean. These smoothed, observed values are based on final, unsmoothed monthly means through June 2005 and on provisional ones thereafter. We compute a smoothed monthly mean by forming the arithmetic average of two sequential 12-month running means of monthly means.

Table entries with numbers in parentheses below them denote predictions by the McNish-Lincoln method. This method estimates future numbers by adding a correction to the mean of all cycles that is proportional to the departure of earlier values of the current cycle from the mean cycle. (See page 9 in the July 1987 supplement to *Solar-Geophysical Data*). We use and predict only smoothed monthly means, because we believe the errors are too great to estimate any values more precise. In the table above,

adding the number in parentheses to the predicted value generates the upper limit of the 90% confidence interval; subtracting the number from the predicted value generates the lower limit. Consider, for example the April 2006 prediction. There exists a 90% chance that in April 2006, the actual smoothed sunspot number will fall somewhere between 10 and 30.

The McNish-Lincoln prediction method generates useful estimates of smoothed, monthly mean sunspot numbers for no more than 12 months ahead. Beyond a year these predictions regress rapidly toward the mean of all 13 cycles used in the computation. Moreover, the method is very sensitive to the date defined as the beginning of the current sunspot cycle, that is, to the date of the most recent sunspot minimum. The new cycle predictions tabulated above are based on the consensus minimum value of 8.8 that occurred in October 1996. For solar maximum discussions, visit <http://www.sec.noaa.gov>.

Although every effort has been made to ensure that these data are correct, we can assume no liability for any damages their inaccuracies might cause. The charge for a 1-year subscription to this monthly bulletin is US\$17.00. To become a subscriber, you may either call (303) 497-6346 or write the NATIONAL GEOPHYSICAL DATA CENTER, Solar-Terrestrial Physics Division (E/GC2), 325 Broadway, Boulder, Colorado 80305-3328 USA. Please include with your written order a cheque or money order payable in U.S. currency to the "Department of Commerce, NOAA/NGDC". Payment may also be made through VISA, MasterCard or American Express credit cards.